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The globular cluster system of M31

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Abstract. This review presents the current status of our knowledge of M31 star clusters. Given the broadness of the subject, I chose to focus on some of its aspects which are not covered by the other participants in this conference.

1. Scientific issues

Globular clusters are studied for their intrinsic properties as well as for the information they give on their parent galaxy. No less than 300 articles address the globular cluster system of M31. I briefly mention here the topics tackled along the years and will come back in more details to some of them below : their total number and specific frequency, their stellar population: age and metallicity, the distance to M31, the relation between the level of the horizontal branch and metallicity, the search for novae, the bulge/disc/halo cluster properties, the radial distribution of their properties, their luminosity function, their kinematics, dynamics, their true shape, ellipticity, tidal radii, effective radii, the absorption law in M31, the mass and mass distribution of M31, the scenario for the formation of M31. For all these topics, comparisons with our Galaxy and galaxies in the Local group are usually conducted.

2. Detection

The clear identification and precise location of the clusters is the first step before any further investigation. There is a very long list of contributors, and it is nearly impossible to quote them all here. Though I would like to mention a few historical cornerstones. Hubble (1932) was the first to report the detection on photographic plates of 140 *nebulous* objects, provisionally identified as globular clusters. Thirty years later, in 1962, Vetesnik reported the existence of 257 candidates. As a matter of fact, the total number of "confirmed" clusters has not changed much over the years. What has changed is the degree of confidence that could be put in their "confirmation". In 1976, Alloin, Pelat & Bijaoui started to focus on the inner regions of M31, and identified 8 central globular clusters. Ten years later, Wirth, Smarr & Bruno (1985) increased the list of bulge clusters up to 30 candidates. In 1979, Hodge initiated the search for open clusters, and proposed a list of 403 such objects. One had to wait till 1991 for the first CCD photometry of 82 halo clusters provided by the work of Racine at the CFHT.

As to the contemporary studies of globular clusters, they are very much based on the major study of Battistini et al. (1993). The authors listed a grand total of 1028 globular cluster candidates, with $14 \text{ mag} < V < 20 \text{ mag}$, in a $3^\circ \times 3^\circ$ region (30 kpc x 30 kpc), centered on M31. They established their *best sample*, which term is now widely used, by selecting clusters with $V \leq 18 \text{ mag}$ and with $B-V \geq 0.58 \text{ mag}$ (this blue limit corresponds to the bluest globular cluster at $B-V = 0.52$ in our Galaxy). This sample regroups 298 objects, comprising 199 confirmed globular clusters, 72 by high resolution imaging, 174 by spectroscopy, 47 by both methods. Limiting the comparison between Galactic and M31 in regions where completeness and uncontamination apply, Battistini et al. (1993) found $N_{M31} / N_{MW} \sim 2.5$. Recently, Barmby et al. (2000) compiled new photometry, spectroscopy and existing data from the literature. They inventory 435 clusters and cluster candidates, including the *best sample*, 268 with 4 or more optical filter photometry, 224 with IR photometry, 200 with velocities, 188 with spectroscopic metallicities. M.G. Lee presents in this conference another new survey.

As to the open clusters, Williams and Hodge (2001) provide an updated catalogue of candidates.

3. A view from the IR

The InfraRed offers a view of the M31 globular cluster system, the closest to the visible.

So far, analyses have dealt mainly with integrated photometry. Cohen & Matthews (1994) gave a good summary of the situation by combining previous works with 23 new globular clusters, 16 of which have projected distance less than 1.1 kpc from M31 center. Among these precedent works, Frogel et al. (1980), Sitko (1984), Bonoli et al. (1992). All works combined led to a total sample of 84 objects. The mean metallicity in M31 and in our Galaxy is the same to within 0.11 dex, which is within the uncertainties of the two $[\text{Fe}/\text{H}]$ determinations. There is no evidence among the outer globular clusters for a spatial gradient of metallicity in the M31 globular cluster sample, as what is observed in our Galaxy.

Stephens et al. present in these proceedings their HST-NICMOS observations of M31 some metal rich globular clusters, a study which opens the era of resolution of individual stars in the IR.

As to the spectroscopy, Davidge (1990) presented spectra in the interval 1.5-1.8 microns and 2.0-2.4 microns of a few luminous M31 clusters. However these spectra being noisy, Davidge had to build a mean spectrum, from which it appeared that the CO bands are similar to the Galactic clusters while CN absorptions look stronger. The CN enhancement in M31 was first mentioned in the visible (Burstein et al. 1984).

4. A view from the UV

Studying the hot phases of the stellar evolution is one strong motivation for observing globular clusters in the UV, another one being the understanding of the connection with the “UV-upturn” population in bulges and elliptical galaxies.

Bohlin et al. (1993) observed with IUT 43 clusters; 42 were identified in the near UV (2490 Å) and only 10 in the far UV (1520 Å). Their data suggest that the M31 clusters are close counterparts of Galactic globulars.

Cacciari et al. (1982) and Cowley & Burstein (1988) obtained the first spectra of a dozen of the brightest clusters in M31, at rather low resolution. Some clusters seem to exhibit residual flux below 3000 Å, greater than that expected from the bright evolved stars in the clusters. Though, these pioneer observations were still at low signal to noise and triggered discussions (Crotts et al. 1990).

The Hubble Space Telescope has opened a new era: the spectral resolution of the UV data can now match that obtained in the visible. As a start, Ponder et al. (1998) got HST/FOS integrated UV spectra for 4 M31 clusters. Interestingly, the near-UV indices involving nitrogen show an enhancement relatively to Galactic stars (disk and halo). As in the visible, there is yet no clear explanation for this phenomenon.

5. A view from the X-rays

It seems that there is a promising field of research in the X-rays. This domain allows a direct insight into the composition of the globular cluster stellar population.

The Einstein Satellite has brought a first set of significant information (Battistini et al. 1982, Trinchieri & Fabbiano 1991). However, ROSAT gave a more general view (Supper et al. 1997, after Primi et al. 1993), since the ROSAT PSPC deep survey could cover the whole galaxy, with a total integration time about ten times greater than that of the Einstein observations. In total, 396 sources were found in the ROSAT PSPC first survey of M31, and 27 of these ROSAT sources were identified with globular clusters. The fraction of X-ray bright clusters among the total cluster population looked similar in M31 and in our Galaxy. For the two galaxies, the maximum luminosity of X-ray bright globular clusters looked also comparable. With the second ROSAT PSPC deep survey, Supper et al. (2001) update and confirm these results. However, *Chandra* results, with observation of 30 clusters, as presented by Di Stefano et al. (2001), seems to contradict this view. Indeed, the authors find a higher peak in luminosity and a larger fraction of high luminosity clusters.

At this stage, it is difficult to distinguish for which reason such opposite conclusions can be reached. Samples, statistical analyses are different. Further investigation seem highly needed.

Very interestingly, the X-ray technology now allows rather good spectroscopy and this first step is reported by Trinchieri et al (1999) with the observation of 8 clusters. The detected sources with BeppoSAX have high X-ray luminosity ($L_X \geq 5 \times 10^{37}$ erg s⁻¹ in the 2-10 keV band). This suggests that they are most likely Low-Mass-X-ray Binaries. Most of them have similar spectrum, that can be described with a single temperature component with $kT \sim 6-9$ keV. Contrary to a recent proposal, no trend shows up between of the soft X-ray properties and the cluster metallicities.

6. A view with the Space Telescope

At the time of this Symposium, about twenty M31 clusters have been imaged by the Space Telescope with enough spatial resolution and photometric quality so that their color-magnitude diagrams could be built. Figure 1 indicates their location in projection over M31. Even if the statistics remain modest, halo, disk and bulge cluster populations have been sampled.

Figure 2 displays a variety of cluster morphologies. Thanks to the courtesy of P. Hodge and B. Williams, the first two upper images are young open and compact clusters, viz. G44 and G38. Below, on the left hand side, thanks to the courtesy of C. Cacciari, L. Federici and F. Fusi Pecci, the disk globular cluster G219. On the right side, G1, the most distant halo cluster. The image is a composite of V and I images (Meylan et al. 2001). This cluster is very bright and massive, and displays a significant flattening. Meylan details some of its outstanding properties in these proceedings. G177, at the bottom of the figure, is a bulge cluster. It is conspicuous that the field surrounding this bulge cluster is significantly more populated than the fields surrounding the other clusters. Due to dynamical interaction with the bulge, G177 has probably lost a significant fraction of the stars in its envelope. Its stellar population is very red, as would be an old and metal rich stellar system.

6.1. An example of color-magnitude diagrams: three bulge star clusters

At the distance of M31, the best color-magnitude diagrams reach stars about one to two magnitudes below the horizontal branch level ($I \sim 24.5$ mag). See, for example, Ajhar et al. (1996), Fusi Pecci et al. (1996), Rich et al. (1996), and Holland et al. (1997). See also B. Williams in these proceedings for the young star clusters.

As seen in Figure 1, most clusters observed with HST are located in the disk and halo of M31. Jablonka et al. (2000) have focussed on three clusters projected at distances between slightly less and slightly more than 1 kpc from the nucleus of M31. The cluster G177 seen in Figure 2 is one of them. Given the extreme density of field stars in the bulge of M31 (from 30 to 55 stars per arcsec² for $V \leq 26.5$ mag and $I \leq 24.5$ mag) and the compactness of the clusters, we adopted the MCS deconvolution technique (Magain, Courbin & Sohy, 1998). The principle of this method is to consider that, in order to satisfy the sampling theorem, the deconvolution must be conducted with a PSF which is narrower than the observed one. One of its appreciable advantages is that it enables the user to perform photometry on the deconvolved images. With this strategy we succeeded in disentangling field from cluster stars, although marginally for the innermost cluster G198. Comparison between field and cluster stars color-magnitude diagrams confirm that these clusters were formed from the same material as the field stars, early in the formation of M31.

6.2. Surface-brightness profiles

The precision achieved with the HST is such that the questions addressed from the surface-brightness profiles resemble more and more those tackled for the clusters in our Galaxy. The structural parameters are derived (Fusi Pecci et

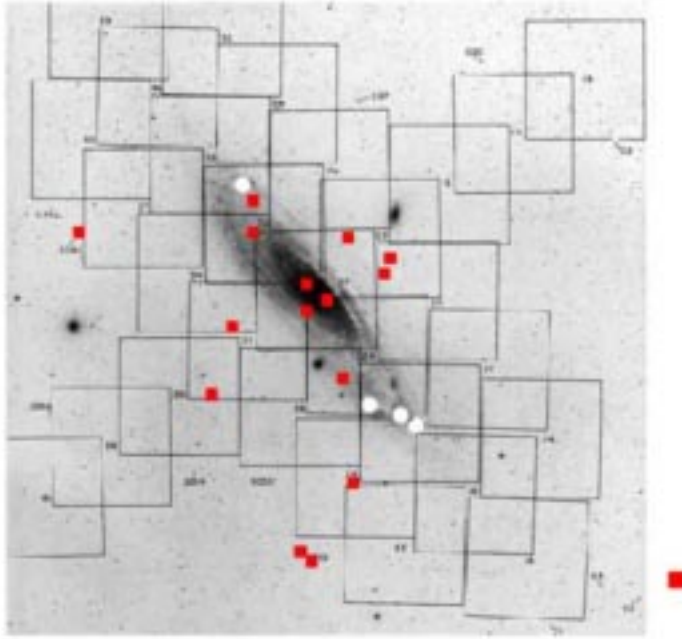


Figure 1. The M31 clusters for which color-magnitude diagrams have been built, at the time of Symposium. Overplotted on the M31 Atlas of Hodge (1961), the squares locate the observed globular clusters, while the circles indicate the position of the observed young clusters

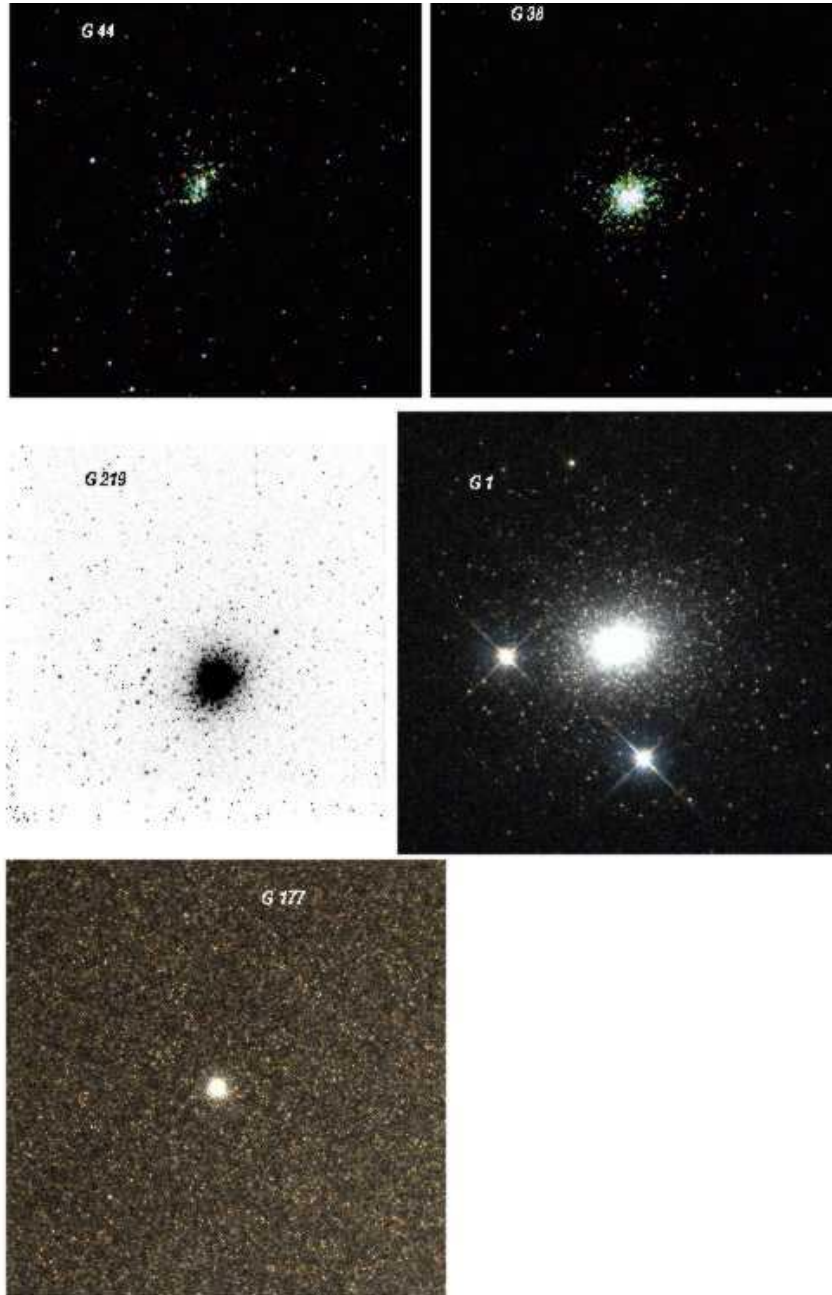


Figure 2. A sample of 5 M31 globular clusters observed with HST/WFPC2: the young clusters G44 and G38, the disk cluster G219, G1 located in the outer halo, and the bulge cluster G177.

al. 1994 ; Grillmair et al. 1996); tidal radii are computed (Cohen & Freeman, 1991); even a collapse core cluster has been identified (Bendinelli et al. 1993).

6.3. Mass-to-Light ratios

Obtaining the velocity dispersion of the globular clusters and combining them with structural and photometric parameters allows to measure the M/L ratios, to explore the correlations of cluster properties and to compare them with the equivalent correlations for the Galactic globular clusters. Two contemporary works started to address this aspect: Dubath & Grillmair (1997) and Djorgovski et al. (1997). The latter, with 27 clusters, certainly gather the largest sample. The authors investigated the correlations in two parameters planes constructed with the central velocity dispersions and central surface brightnesses or absolute magnitudes, both in V and K bands. They stress that the M31 globular clusters smoothly continue and extend to higher luminosities the trends drawn by the Galactic clusters. Moreover, at same metallicity, Galactic and M31 clusters have the same M/L ratios.

The so far most precise dynamical study of a globular cluster in M31, viz. G1, based on its surface-brightness profile and central velocity dispersion, has been recently published by Meylan et al. (2001 and these proceedings).

7. Conclusion

It clearly appears that one will gain in increasing the size of the cluster samples for which the quality of the data will catch up with what is available for the clusters in our Galaxy, as one started with the HST. The answer to the question of the universality of the cluster properties is at this price. Meanwhile, the opening of new wavelength domains with good sensitivity and improvement of spatial/spectral resolutions certainly give to the study of globular clusters in M31 some fresh air ! This is becoming a new domain of investigation in which excitement for new insight is genuine.

References

- Ajhar E.A., Grillmair C.J., Lauer T.R. et al., 1996, AJ, 111, 1110
- Barmby P., Huchra J.P., Brodie J.P. et al. 2000, ApJ, 119, 247
- Battistini P.L., Bonoli F., Pecci F.F. et al., 1982 A&A 113 39
- Battistini P.L., Bonoli F., Casavecchia M. et al. 1993, A&A, 272
- Bendinelli O., Cacciari C., Djorgovski S. et al., 1993, ApJL, 409
- Bohlin R.C., Deutsch E.W., McQuade K.A. et al. 1993, ApJ, 417, 127
- Bohlin R.C., Cornett R.H., Hill J.K. et al. 1988, ApJ, 334, 657
- Bonoli F., Delpino F., Federici L. et al. 1992, A&AS, 96, 163
- Burstein D., Faber S.M., Gaskell C.M., Krumm N., 1984, ApJ, 287, 586
- Cacciari C., Cassatella A., Bianchi L. et al., 1982, ApJ, 261, 77
- Cohen J.G., Freeman K.C., 1991, AJ, 101, 483
- Cohen J.G., Matthews K., 1994, AJ, 108, 128

- Cowley A.P., Burstein D., 1988, AJ, 95, 1071
Crotts A.P.S., Kron R.G., Cacciari C., Fusi Pecci F., 1990, AJ, 100, 141
Davidge T.J., 1990, ApJL, 351, 37
Di Stefano R., Kong A.K.H., Garcia M.R., Barmby P., 2001, astro-ph/0106254
Djorgovski S.G., Gal R.R., McCarthy J.K. et al., 1997, ApJL, 474, 19
Dubath P., Grillmair C.J., 1997, A&A, 321, 379
Frogel J.A., Persson S.E., Cohen J.G., 1980, ApJ, 240, 785
Fusi Pecci F., Battistini P., Bendinelli O. et al., 1994 A&A 284 349
Fusi Pecci F., Buonanno R., Cacciari C. et al., 1996, AJ, 112, 1461
Grillmair C.J., Ajhar E.A., Faber S.M. et al., 1996 AJ 111 2293
Holland S., Fahlman G.G., Richer H.B., 1997, AJ, 114, 1488
Jablonka P., Courbin F., Meylan G. et al., 2000, A&A, 359, 131
Magain P., Courbin F., Sohy S., 1998, ApJ, 494, 472
Ponder J.M., Burstein D., O'Connell R.W. et al., 1998, AJ, 116, 2297
Primini F.A., Forman W., Jones C., 1993, ApJ, 410, 615
Racine R., 1991, AJ, 101, 865
Rich R.M., Mighell K.J., Freedman W.L., Neill J.D., 1996 AJ 111 768
Sitko, M.L., 1984, ApJ, 286, 209
Supper R., Hasinger G., Pietsch W. et al., 1996, AJ, 112, 1461
Supper R., Hasinger G., Lewin W.H.G. et al., 2001, A&A 373, 63-99
Trinchieri G., Fabbiano G., 1991, ApJ, 382, 82
Trinchieri G., Israel G.L., Chiappetti L., et al. 1999, A&A, 348, 43
Vetesnik M., 1962, BAICz, 13, 180
Williams B.F., Hodge P.W., 2001, astro-ph/015506